## Effect of intermediate Minkowskian evolution on CMB bispectrum

S. Mironov

INR RAS & ITEP

Suzdal, June, 7th, 2014

based on arXiv:1312.7808 (S.M., S.Ramazanov, V.Rubakov)

## • Perturbations in the Universe are approximately Gaussian and have flat spectrum

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field
- Flatness may be due to some symmetry

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field
- Flatness may be due to some symmetry
- There should be some early-Universe model

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field
- Flatness may be due to some symmetry
- There should be some early-Universe model
- inflation (many types)

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field
- Flatness may be due to some symmetry
- There should be some early-Universe model
- inflation (many types)
- conformal rolling, galilean Genesis

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field
- Flatness may be due to some symmetry
- There should be some early-Universe model
- inflation (many types)
- conformal rolling, galilean Genesis
- ekpyrosis, starting, bouncing

- Perturbations in the Universe are approximately Gaussian and have flat spectrum
- First property suggests they originate from amplified vacuum fluctuations of weakly coupled quantum field
- Flatness may be due to some symmetry
- There should be some early-Universe model
- inflation (many types)
- conformal rolling, galilean Genesis
- ekpyrosis, starting, bouncing
- every model has specific features (tilt, statistical anisotropy, non-gaussianity)

• We are studying how any of this non-gaussianity is changed due to an additional stage

- We are studying how any of this non-gaussianity is changed due to an additional stage
- Conformal rolling example:

- We are studying how any of this non-gaussianity is changed due to an additional stage
- Conformal rolling example:



- We are studying how any of this non-gaussianity is changed due to an additional stage
- Conformal rolling example:



by the end of the rolling, perturbations of interest already have flat spectrum (and special form of non-Gaussianity)

- they are superhorizon and remain until the hot epoch

 they are superhorizon and remain until the hot epoch thus perturbation do not evolve, form of non-Gaussianity is not changed

- they are superhorizon and remain until the hot epoch thus perturbation do not evolve, form of non-Gaussianity is not changed
- they have a period of subhorizon evolution before the hot stage

- they are superhorizon and remain until the hot epoch thus perturbation do not evolve, form of non-Gaussianity is not changed
- they have a period of subhorizon evolution before the hot stage perturbations oscillate in nearly *Minkowski* space-time

- they are superhorizon and remain until the hot epoch thus perturbation do not evolve, form of non-Gaussianity is not changed
- they have a period of subhorizon evolution before the hot stage perturbations oscillate in nearly *Minkowski* space-time [otherwise spectrum will be grossly modified]

- they are superhorizon and remain until the hot epoch thus perturbation do not evolve, form of non-Gaussianity is not changed
- they have a period of subhorizon evolution before the hot stage perturbations oscillate in nearly *Minkowski* space-time [otherwise spectrum will be grossly modified]
- we consider the second situation

we assume perturbations to be massless and non-interacting during the Minkowski stage

$$\delta heta(\mathbf{k},\eta) = \delta heta(\mathbf{k},\eta_*) \cos k(\eta_*-\eta)$$

we assume perturbations to be massless and non-interacting during the Minkowski stage

$$\delta \theta(\mathbf{k},\eta) = \delta \theta(\mathbf{k},\eta_*) \cos k(\eta_* - \eta)$$

or after horizon exit (time of the exit  $\eta_{ex} = 0$ )

$$\delta\theta(\mathbf{k}) = \delta\theta(\mathbf{k},\eta_*)\cos k\eta_*$$

we assume perturbations to be massless and non-interacting during the Minkowski stage

$$\delta\theta(\mathbf{k},\eta) = \delta\theta(\mathbf{k},\eta_*)\cos k(\eta_*-\eta)$$

or after horizon exit (time of the exit  $\eta_{ex} = 0$ )

$$\delta\theta(\mathbf{k}) = \delta\theta(\mathbf{k},\eta_*)\cos k\eta_*$$

let us present a correlator in terms of the primordial Newtonian potential

$$\langle \Phi(\mathbf{k}_1) \Phi(\mathbf{k}_2) \Phi(\mathbf{k}_3) \rangle =$$
  
=  $\left(\frac{\pi \mathcal{P}_{\Phi}}{2}\right)^{3/2} A(k_1, k_2, k_3) \delta\left(\sum_i \mathbf{k}_i\right) \cos(k_1 \eta_*) \cos(k_2 \eta_*) \cos(k_3 \eta_*)$ 

 $A(k_1, k_2, k_3)$  is initial shape function

$$a_{lm} = \int d\mathbf{n} \delta T(\mathbf{n}) Y^*_{lm}(\mathbf{n})$$

・ロト ・聞ト ・ヨト ・ヨト

Ξ.

$$a_{lm} = \int d\mathbf{n} \delta T(\mathbf{n}) Y_{lm}^*(\mathbf{n})$$

$$a_{lm} = 4\pi i^{l} \int \frac{d\mathbf{k}}{(2\pi)^{3/2}} \Delta_{l}(k\eta_{0}) \Phi(\mathbf{k}) Y_{lm}^{*}(\hat{\mathbf{k}})$$

$$a_{lm} = \int d\mathbf{n} \delta T(\mathbf{n}) Y_{lm}^*(\mathbf{n})$$

$$a_{lm} = 4\pi i^{l} \int \frac{d\mathbf{k}}{(2\pi)^{3/2}} \Delta_{l}(k\eta_{0}) \Phi(\mathbf{k}) Y_{lm}^{*}(\hat{\mathbf{k}})$$

- Recall that  $\Delta_I(y) \propto j_I(y)$ , where  $j_I$  is the spherical Bessel function

$$a_{lm} = \int d\mathbf{n} \delta T(\mathbf{n}) Y_{lm}^*(\mathbf{n})$$

$$a_{lm} = 4\pi i^{l} \int \frac{d\mathbf{k}}{(2\pi)^{3/2}} \Delta_{l}(k\eta_{0}) \Phi(\mathbf{k}) Y_{lm}^{*}(\hat{\mathbf{k}})$$

- Recall that  $\Delta_I(y) \propto j_I(y)$ , where  $j_I$  is the spherical Bessel function let us denote:  $\mathbf{y}_i \equiv \mathbf{k}_i \eta_0$ ,  $z \equiv -\frac{\eta_*}{\eta_0}$ ,  $A(y_1, y_2, y_3) \equiv \frac{A(k_1, k_2, k_3)}{\eta_0^6}$ 

$$a_{lm} = \int d\mathbf{n} \delta T(\mathbf{n}) Y_{lm}^*(\mathbf{n})$$

$$a_{lm} = 4\pi i^{l} \int \frac{d\mathbf{k}}{(2\pi)^{3/2}} \Delta_{l}(k\eta_{0}) \Phi(\mathbf{k}) Y_{lm}^{*}(\hat{\mathbf{k}})$$

- Recall that  $\Delta_l(y) \propto j_l(y)$ , where  $j_l$  is the spherical Bessel function let us denote:  $\mathbf{y}_i \equiv \mathbf{k}_i \eta_0$ ,  $z \equiv -\frac{\eta_*}{\eta_0}$ ,  $A(y_1, y_2, y_3) \equiv \frac{A(k_1, k_2, k_3)}{\eta_0^6}$ we assume the ordering  $l_1 \leq l_2 \leq l_3$  in what follows

$$\langle a_{l_1 m_1} a_{l_2 m_2} a_{l_3 m_3} \rangle = i^{l_1 + l_2 + l_3} \mathcal{P}_{\Phi}^{3/2} \int d\mathbf{y}_1 d\mathbf{y}_2 \Delta_{l_1}(y_1) \Delta_{l_2}(y_2) \Delta_{l_3}(|\mathbf{y}_1 + \mathbf{y}_2|) \times \\ \times Y_{l_1 m_1}^*(\theta_1, \phi_1) Y_{l_2 m_2}^*(\theta_2, \phi_2) Y_{l_3 m_3}^*(\theta_3, \phi_3) \times \\ \times \cos(y_1 z) \cos(y_2 z) \cos(|\mathbf{y}_1 + \mathbf{y}_2| z) A(y_1, y_2, y_3)$$

・ロト ・四ト ・ヨト ・ヨト

2

$$\langle a_{l_1 m_1} a_{l_2 m_2} a_{l_3 m_3} \rangle = i^{l_1 + l_2 + l_3} \mathcal{P}_{\Phi}^{3/2} \int d\mathbf{y}_1 d\mathbf{y}_2 \Delta_{l_1}(y_1) \Delta_{l_2}(y_2) \Delta_{l_3}(|\mathbf{y}_1 + \mathbf{y}_2|) \times \\ \times Y_{l_1 m_1}^*(\theta_1, \phi_1) Y_{l_2 m_2}^*(\theta_2, \phi_2) Y_{l_3 m_3}^*(\theta_3, \phi_3) \times \\ \times \cos(y_1 z) \cos(y_2 z) \cos(|\mathbf{y}_1 + \mathbf{y}_2| z) A(y_1, y_2, y_3)$$

we discuss two regimes.

3

$$\langle a_{l_1m_1}a_{l_2m_2}a_{l_3m_3}\rangle = i^{l_1+l_2+l_3} \mathcal{P}_{\Phi}^{3/2} \int d\mathbf{y}_1 d\mathbf{y}_2 \Delta_{l_1}(y_1) \Delta_{l_2}(y_2) \Delta_{l_3}(|\mathbf{y}_1+\mathbf{y}_2|) \times \\ \times Y_{l_1m_1}^*(\theta_1,\phi_1) Y_{l_2m_2}^*(\theta_2,\phi_2) Y_{l_3m_3}^*(\theta_3,\phi_3) \times \\ \times \cos(y_1z) \cos(y_2z) \cos(|\mathbf{y}_1+\mathbf{y}_2|z) A(y_1,y_2,y_3)$$

we discuss two regimes.

- first one: saddle point contribution directly from cosine (spherical functions are assumed to be slow functions)

$$\langle a_{l_1m_1}a_{l_2m_2}a_{l_3m_3}\rangle = i^{l_1+l_2+l_3} \mathcal{P}_{\Phi}^{3/2} \int d\mathbf{y}_1 d\mathbf{y}_2 \Delta_{l_1}(y_1) \Delta_{l_2}(y_2) \Delta_{l_3}(|\mathbf{y}_1+\mathbf{y}_2|) \times \\ \times Y_{l_1m_1}^*(\theta_1,\phi_1) Y_{l_2m_2}^*(\theta_2,\phi_2) Y_{l_3m_3}^*(\theta_3,\phi_3) \times \\ \times \cos(y_1z) \cos(y_2z) \cos(|\mathbf{y}_1+\mathbf{y}_2|z) A(y_1,y_2,y_3)$$

we discuss two regimes.

- first one: saddle point contribution directly from cosine (spherical functions are assumed to be slow functions)
- second one is more general: we use approximation for spherical functions

$$Y_{lm}(\theta,\phi) = \frac{1}{\pi\sqrt{\sin\theta}} \cos\left[\left(l+\frac{1}{2}\right)\theta - \frac{\pi}{4} + \frac{\pi m}{2}\right] e^{im\phi} + \mathcal{O}\left(\frac{1}{l}\right)$$

7 / 10

In general case the answer is

$$\begin{split} \langle a_{l_1 m_1} a_{l_2 m_2} a_{l_3 m_3} \rangle &= i^{l_1 + l_2 + l_3} \cdot \frac{\pi \mathcal{P}_{\Phi}^{3/2}}{4} \cdot \mathcal{B}_{l_2 m_2; l_3 m_3}^{l_1 m_1} \times \\ \times \int dy_1 dy_2 \cdot y_1^2 \cdot y_2^2 \cdot \Delta_{l_1}(y_1) \Delta_{l_2}(y_2) \Delta_{l_3}(y_1 + y_2) \cdot \mathcal{A}(y_1, y_2, y_1 + y_2) \times \\ & \times \frac{2(y_1 + y_2)}{zy_1 y_2} \cdot \sin\left\{\frac{1}{2} \cdot \frac{y_2^2 l_1^2 + y_1^2 l_2^2}{y_1 y_2(y_1 + y_2) z}\right\} \cdot \cos\left\{\frac{l_1 l_2}{(y_1 + y_2) z}\right\} \times \\ & \times \left[1 - \cot\left\{\frac{1}{2} \cdot \frac{y_2^2 l_1^2 + y_1^2 l_2^2}{y_1 y_2(y_1 + y_2) z}\right\} \cdot \tan\left\{\frac{l_1 l_2}{(y_1 + y_2) z}\right\} \cdot \frac{\mathcal{B}_{l_2, -1; l_3 0}^{l_1, 1}}{\mathcal{B}_{l_2 0; l_3 0}^{l_1, 0}}\right] \\ & + (l_1, m_1 \leftrightarrow l_3, m_3) + (l_2, m_2 \leftrightarrow l_3, m_3) \end{split}$$

э

In general case the answer is

$$\begin{split} \langle a_{l_1 m_1} a_{l_2 m_2} a_{l_3 m_3} \rangle &= i^{l_1 + l_2 + l_3} \cdot \frac{\pi \mathcal{P}_{\Phi}^{3/2}}{4} \cdot \mathcal{B}_{l_2 m_2; l_3 m_3}^{l_1 m_1} \times \\ \times \int dy_1 dy_2 \cdot y_1^2 \cdot y_2^2 \cdot \Delta_{l_1}(y_1) \Delta_{l_2}(y_2) \Delta_{l_3}(y_1 + y_2) \cdot \mathcal{A}(y_1, y_2, y_1 + y_2) \times \\ & \times \frac{2(y_1 + y_2)}{zy_1 y_2} \cdot \sin\left\{\frac{1}{2} \cdot \frac{y_2^2 l_1^2 + y_1^2 l_2^2}{y_1 y_2(y_1 + y_2) z}\right\} \cdot \cos\left\{\frac{l_1 l_2}{(y_1 + y_2) z}\right\} \times \\ & \times \left[1 - \cot\left\{\frac{1}{2} \cdot \frac{y_2^2 l_1^2 + y_1^2 l_2^2}{y_1 y_2(y_1 + y_2) z}\right\} \cdot \tan\left\{\frac{l_1 l_2}{(y_1 + y_2) z}\right\} \cdot \frac{\mathcal{B}_{l_2, -1; l_3 0}^{l_1, 1}}{\mathcal{B}_{l_2 0; l_3 0}^{l_1, 0}}\right] \\ & + (l_1, m_1 \leftrightarrow l_3, m_3) + (l_2, m_2 \leftrightarrow l_3, m_3) \end{split}$$

let us proceed to the qualitative results

The bispectrum vanishes for  $l_1 + l_2 < l_3$  and undergoes oscillations as function of  $\Delta l = l_1 + l_2 - l_3 > 0$  with roughly constant amplitude.

The bispectrum vanishes for  $l_1 + l_2 < l_3$  and undergoes oscillations as function of  $\Delta l = l_1 + l_2 - l_3 > 0$  with roughly constant amplitude. [The origin of oscillations can be traced back to oscillations in the CMB transfer functions]

The bispectrum vanishes for  $l_1 + l_2 < l_3$  and undergoes oscillations as function of  $\Delta l = l_1 + l_2 - l_3 > 0$  with roughly constant amplitude. [The origin of oscillations can be traced back to oscillations in the CMB transfer functions]

the fact, that amplitude is approximately constant is quite non-trivial and important for discriminating the models we discussed in this paper from inflationary scenarios, some of which predict the oscillatory behaviour of the bispectrum analogous to ours, but strongly peaked in flattened triangle limit.

## THANK YOU FOR YOUR ATTENTION!

3

-