## Particle production in Starobinsky model of dark energy

#### A.A. Tokareva in collaboration with D.S. Gorbunov

MSU, INR

June, 5

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★  $M = 3 \times 10^{13}$  GeV is appropriate for the Starobinsky inflation

A. A. Starobinsky, "Disappearing cosmological constant in f(R) gravity," JETP Lett. 86, 157 (2007) [arXiv:0706.2041 [astro-ph]]

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$$V(\phi) = \frac{M_P^2}{2F'(R(\phi))^2} \left(RF'(R) - F(R)\right),$$

$$F'(R) = e^{2\phi/(\sqrt{6}M_P)} \rightarrow R(\phi)$$

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ho - 3p}{4}e^{-4ar{\phi}}, \ \ ar{\phi} = rac{\phi}{\sqrt{6}M_P}$$

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In the considering model:

 $\lambda \gg 1$ 

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The scalaron mass:

$$m^2 = V_{eff}''(\phi_{min})$$

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For  $m \ll M$   $m_{eff} \sim \tau^{n+1} \Rightarrow$  particle production when  $\tau$  is changing

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★ In a paper E. V. Arbuzova, A. D. Dolgov and L. Reverberi, "Particle Production in f(R) Gravity during Structure Formation," Phys. Rev. D 88, no. 2, 024035 (2013) [arXiv:1305.5668 [gr-qc]] the calculation of particle production was carried in the Jordan frame. But the choice of initial conditions – pure GR solution – look questionable. We calculate the number of created particles starting from more realistic vacuum initial conditions.

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$$\rho(t) = 
ho_0 e^{t/t_J}$$

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$$\ddot{\phi} + (k^2 + m_0^2 e^{2\beta t})\phi = 0$$

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• Vacuum initial conditions at t = 0:

$$\phi = 1/\sqrt{2\omega}, \ \dot{\phi} = -i\omega\phi, \ \omega = \sqrt{k^2 + m_0^2 e^{2eta t}}$$

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$$n_p = 8.7 \cdot 10^{-87} \mathrm{cm}^{-3} \cdot \frac{(n+1)^2}{\sqrt{12n(2n+1)}} \left(\frac{\lambda}{2}\right)^n \alpha^{n+2}$$

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- contracting stars, supernovae etc. are too dense: scalaron mass is constant
- star formation in our galaxy



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$$\frac{\rho_0}{\rho_{\Lambda}} = \alpha = 3.7 \cdot 10^{16} \, \left(\frac{M_{\odot}}{M_*}\right)^2 \left(\frac{T}{100K}\right)^3.$$

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The spectrum has a cutoff at  $m \sim M$ 

#### Products of scalaron decays: UHE protons



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Proton mean free path in the contracting protostar with radius R:

$$\frac{\lambda_{p}}{R} = 0.05 \frac{M_{*}}{M_{\odot}} \left(\frac{T}{100 \, K}\right)^{-2}$$

#### Estimation of proton flux

$$\frac{dN}{dE} = \kappa N_b \frac{n_p R^3}{r^2 t_J} N \xi(M_*) \frac{dM_*}{dE}$$

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- $\kappa \sim 0.1$  fraction of protons in scalaron decay products
- $N_b \sim 300$  full number of produced particles
- n<sub>p</sub> scalaron number density
- R star radius
- $r \sim 10 \text{ kpc}$  average distance in our galaxy
- t<sub>j</sub> the time of exponential contraction (free fall)
- N number of objects
- $\xi(M_*) = 0.08 (M_*/M_{\odot})^{-2.3}/M_{\odot}$  initial mass function (Salpeter)

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- $\xi(M_*) = 0.08 (M_*/M_{\odot})^{-2.3}/M_{\odot}$  initial mass function (Salpeter)

$$\frac{dN}{dE} \sim E^{\frac{0.15}{n+1}} T^{-1.95} \left(\frac{\lambda}{2}\right)^{-\frac{0.15n}{n+1}}$$

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## UHE CR spectrum



A.A. Tokareva in collaboration with D.S. Gorbunov Particle production in Starobinsky model of dark energy

#### Dependence on temperature, $\lambda$ , *n*





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#### Dependence on temperature, $\lambda$ , *n*

E=3.10<sup>20</sup> eV

3

0.5



n

6

5

3

4

T=10 K E=3.10<sup>20</sup> eV

5

6

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- A calculation of the scalaron creation in contracting objects based on vacuum initial conditions was carried out.
- The proton flux from the star formation processes in our galaxy may give a significant impact to the UHE CR quadrupole anisotropy.
- For some values of the parameter  $n \ (n = 3 \div 6)$  we can conclude that maximal scalaron mass M must be  $M < 9 \times 10^{13}$  GeV.

# Thanks for your attention!

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